

# Spectrophotometry of speckle binary stars

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**Abstract.** Spectrophotometric observations of 20 speckle interferometric binary stars in the wavelength range 3640–8340 Å with a resolution of 18 Å are presented in the form of the measured flux corrected for atmospheric extinction in the units of erg/cm<sup>2</sup> · s·Å, with no corrections for interstellar reddening. *B*, *V*, and *R* magnitudes, *B* – *V* colour indices and entire spectral types are also presented and compared with those of Hipparcos and Tycho catalogues and SIMBAD.

**Key words:** stars: spectrophotometry: spectral energy distribution - stars: binaries: speckle binary stars

## 1. Introduction

The study of binary and multiple systems by means of speckle interferometry made a valuable contribution to the understanding of formation and evolution of stellar systems, especially in the recent years with the aid of large telescopes and utilization of diffraction limited techniques. The direct results of speckle interferometric observations are separation angle, orientation angle, and magnitude difference for the sub-components of each binary or multiple system (Balega et al., 2002). In its turn this leads to the determination of orbit and orbital period. Using other kinds of observations, like high resolution spectroscopy (Tokovinin, 1995) or wide range spectrophotometry (Al-Wardat et al., 2002; Al-Wardat, 2002), the number of deduced parameters can be raised, and wide understanding of such systems can be achieved.

In this paper, a wide range (3640–8340 Å), low resolution (18 Å, 6 Å/px) spectrophotometry of 20 speckle interferometric binary stars is presented. The objects of the study were taken from the speckle interferometric programme, which has been carried out at the 6 m telescope of the Special Astrophysical Observatory since the early 90s. The programme mainly includes late type dwarfs in the vicinity of the Sun, and their fundamental parameters are badly known. The presented data can be used as a reference for building theoretical spectral energy distribution curves on the basis of Kurucz blanketed models. This, along with the magnitude difference from speckle interferometric observations, can be used to build a spectral energy distribution for each of the components from which we can get their  $T_{eff}$ , lg  $g$ , and spectral types.

The stars are listed in Table 1 with different identifications: Hipparcos (Col. 1), HD (Col. 2), other

identifications (Col. 3). The coordinates of the stars (Table 1, Col. 4, 5) were taken from SIMBAD astronomical database.

## 2. Observations and data analysis

Spectra were obtained using a low resolution grating (325/4° grooves/mm, 5.97 Å/px reciprocal dispersion) within the UAGS spectrograph at the Cassegrain focus of the Carl Zeiss Jena (Zeiss-1000) 1 m telescope of SAO during the photometrical nights, January 28 and February 4, 2002. The seeing was around 1.5''. The times of observations in terms of Julian Dates are listed in Table 1, Col.6. The stars AD Leo, Wolf424, and Hip78864, which exhibit variability, were observed at a specific time.

The spectrograph has an ISD015 A 530×580 px CCD detector. A 0.5 mm slit width was used to encompass all light from the star, and it was rotated (by changing the angle of the instrument's table) to a suitable direction to prevent the effect of nearby stars.

Two positional angles for the grating were used to cover the spectral range between 3640 Å and 8340 Å, 29° for the blue part and 30°40' for the red part. This was done for two reasons: first, because of small dimensions of the detector which does not cover this spectral range, and, second, in order to overcome the falling sensitivity of the detector in the blue part of the spectrum by applying longer integration times in this part, and there were at least 500 Å of overlap between the two regions (see Fig. 1), allowing us further checks on our internal agreement.

Standards from Massey et al. (1988), Oke (1990) and Hamuy et al. (1992, 1994) were used for the cal-

ibration of the system. The spectra were sky subtracted and wavelength calibrated. Then the spectral sensitivity curve of the CCD derived from the standard stars' spectra for each angle, was used for flux calibration of the object spectra. All steps were made using ESO-MIDAS<sup>1</sup> routines. The wavelength calibration was performed by means of He-Ne-Ar lamp emission spectra.

For each star, the spectra were averaged separately for angles of the grating. Then by averaging the resulting individual spectra in the overlap region, we obtained a single 4700 Å band spectrum.

As an example, four individual spectra of Hip12552 obtained on February 4th, 2002 are plotted together in Fig. 1. No adjustment has been made in their flux level. The coincidence reflects a good internal agreement of our results.

The standard deviation of *B* and *V* magnitudes, obtained for each star from the sample of the spectra, is typically better than 0<sup>m</sup>06, and for the *R* band it is better than 0<sup>m</sup>07. The error bars are the lowest in the central part of the spectrum where the blue and red spectra overlap.

To investigate the external agreement of our results, we used some of the standard stars. We observed these standards as objects, then we compared the results with the published spectral energy distributions of these standards. We observed the stars HR718 (4<sup>m</sup>28, B9III), HR5501 (5<sup>m</sup>68, B9.5V), and HR1544 (4<sup>m</sup>36, AIV), where Hilt600 (10<sup>m</sup>44, B1) from Massey et al. (1988), Feige66 (10<sup>m</sup>50, sdO) from Massey et al. (1988), and HD93521 (7<sup>m</sup>04, O9Vp) from Oke (1990) were used as standards in obtaining the spectra, respectively. Then we compared our results with those published by Hamuy et al. (1992, 1994) (Figs. 2, 3, and 4). The comparisons show a good agreement within the estimated error values from the internal agreement.

### 3. Results and discussion

Table 2 lists the measured flux of the stars, corrected for the atmospheric extinction, in units of erg/cm<sup>2</sup>·s·Å, where Hipparcos identification, if exists, was used for the stars, else other identifications were used. The data are plotted in Fig. 5. The low sensitivity of the detector in the blue part is noteworthy here, which is clear in the spectra.

*BVR* synthetic magnitudes were computed using the following integrals:

$$X = -2.5 \log \frac{\int S_x(\lambda) F_\lambda d\lambda}{\int S_x(\lambda) d\lambda} + ZP,$$

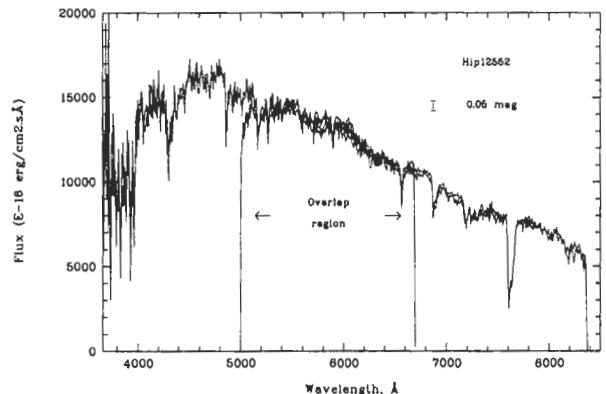


Figure 1: Superposition of 4 different spectra of Hip12552 is shown as an evidence for the final agreement of spectra. The overlap region between blue and red spectra is shown also.

where  $S_x(\lambda)$  is the transmission function for the passband X. We adopted the filter functions  $B_{90}$ ,  $V_{90}$  and  $R_{90}$  published by Bessel (1990).  $ZP$  is the zero point for the magnitude scale. For *V* band  $ZP$  was solved using the spectrophotometric calibration of Vega published by Hayes (1985) and the *V* magnitude of 0<sup>m</sup>03 measured by Johnson et al. (1966). While for *B* and *R* bands it was solved using the Vega magnitudes published by Hamuy et al. (2001) as  $B = 0^m014$ , and  $R = 0^m042$ , since they are more reliable than those obtained by Johnson et al. (1966) (see Appendix B in Hamuy et al., 2001). The integrals were computed after interpolating  $S_x(\lambda)$  to the wavelength spacing of  $F_\lambda^{Star}$  which is 6 Å. The results of these calculations are listed in Table 1, Col. 7, 8, 9, and 10.

Figs. 6, 7, and 8 show comparisons between the calculated *B* magnitudes, *V* magnitudes and *B* – *V* colour indices with Johnson *B*, *V*, and *B* – *V* of Hipparcos catalogue (fields H5 and H37). The Hipparcos magnitudes were taken either from the ground based observations or calculated from  $B_T$  and  $V_T$  of Tycho using different relations for different kinds or luminosity classes of the stars (for more information see Hipparcos and Tycho catalogues sec. 1.3 (ESA, 1997)).

The entire spectral types of the binaries were estimated by comparing *B* – *V* with the intrinsic colours of FitzGerald (1970) neglecting interstellar reddening since all of the stars are nearby stars and their interstellar reddening lies within the error values of *B* – *V* (see Al-Wardat, 2002). Results are listed in Table 1, Col. 11, along with those from SIMBAD (Col. 12) for comparison sake, where they show a good agreement (within the error values of *B* – *V*) for 15 stars, while the other 5 stars: (Hip1055, Hip51945, Hip64838, Hip70973, and Hip78864) show differences between the estimated spectral types in this work and those given by SIMBAD.

<sup>1</sup> Munich Image Data Analysis System, developed, maintained and distributed by the European Southern Observatory.

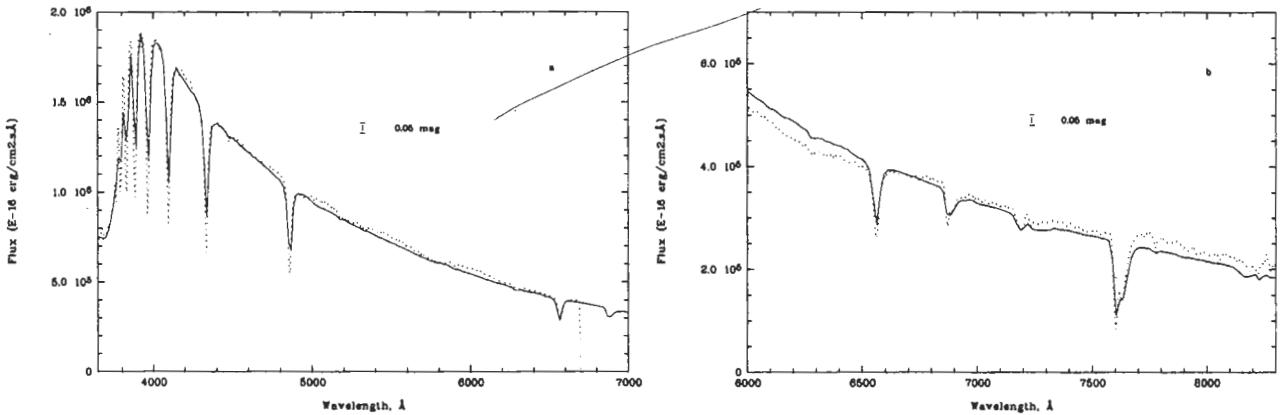


Figure 2: Blue part (a) and red part (b) comparisons between the obtained spectral energy distribution of HR718 ( $4^m 28$ , B9III) (dotted line) and that published by Hamuy et al. (1992, 1994) (solid line). Hilt600 ( $10^m 44$ , B1) from Massey et al. (1988) was used as a standard in obtaining the spectrum.

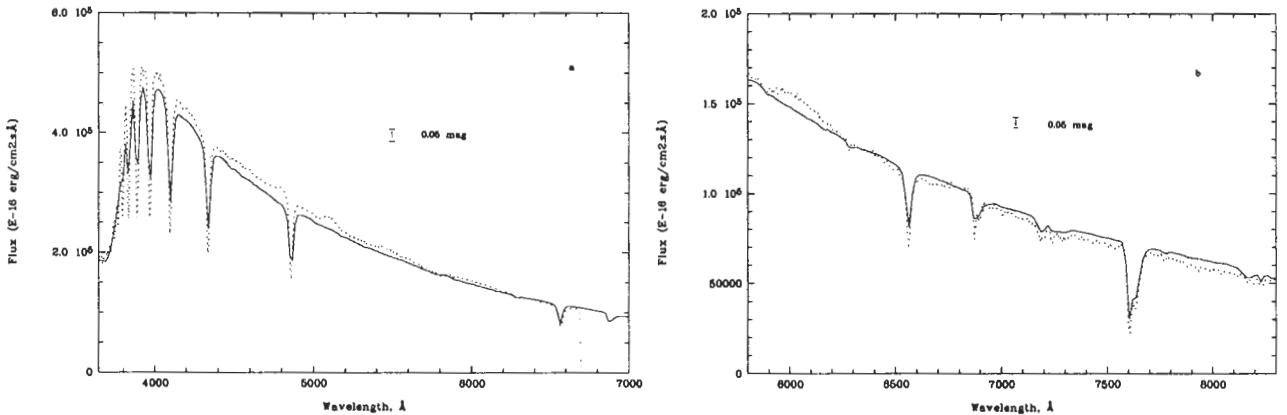


Figure 3: Blue part (a) and red part (b) comparisons between the obtained spectral energy distribution of HR5501 ( $5^m 68$ , B9.5V) (dotted line) and that published by Hamuy et al. (1992, 1994) (solid line). Feige66 ( $10^m 50$ , sdO) from Massey et al. (1988) was used as a standard in obtaining the spectrum.

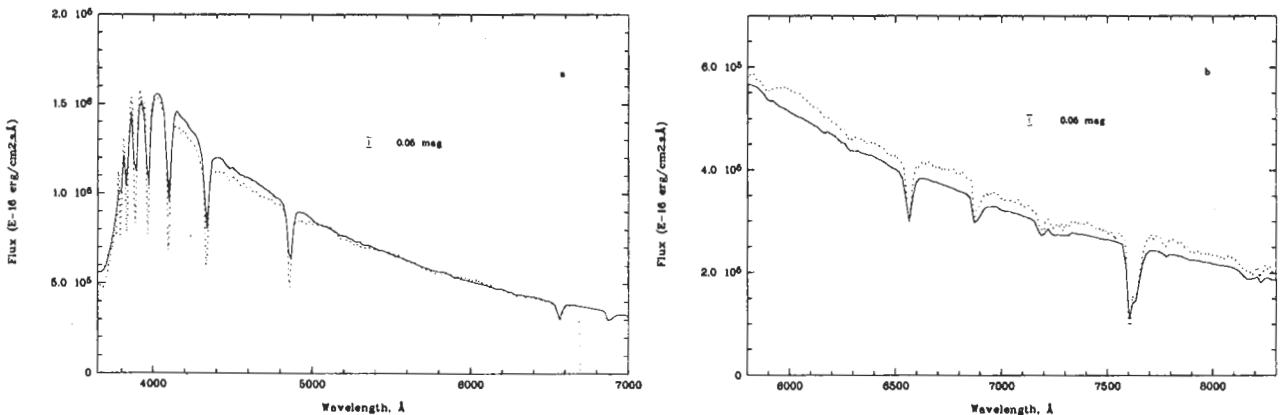


Figure 4: Blue part (a) and red part (b) comparisons between the obtained spectral energy distribution of HR1544 ( $4^m 36$ , AIV) (dotted line) and that published by Hamuy et al. (1992, 1994) (solid line). HD93521 ( $7^m 04$ , O9Vp) from Oke (1990) was used as a standard in obtaining the spectrum.

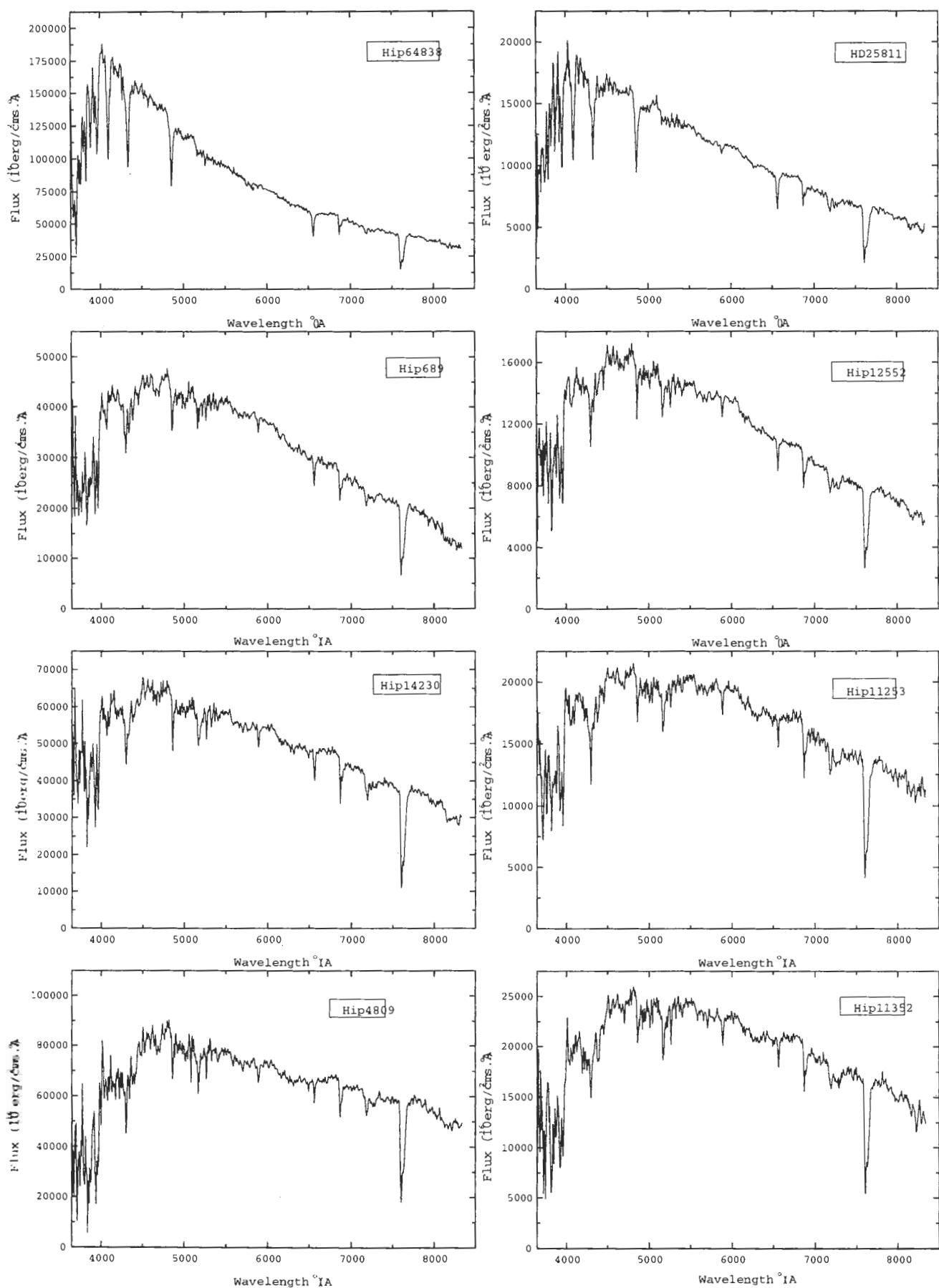


Figure 5: Spectral energy distributions of the stars labeled with the star names.

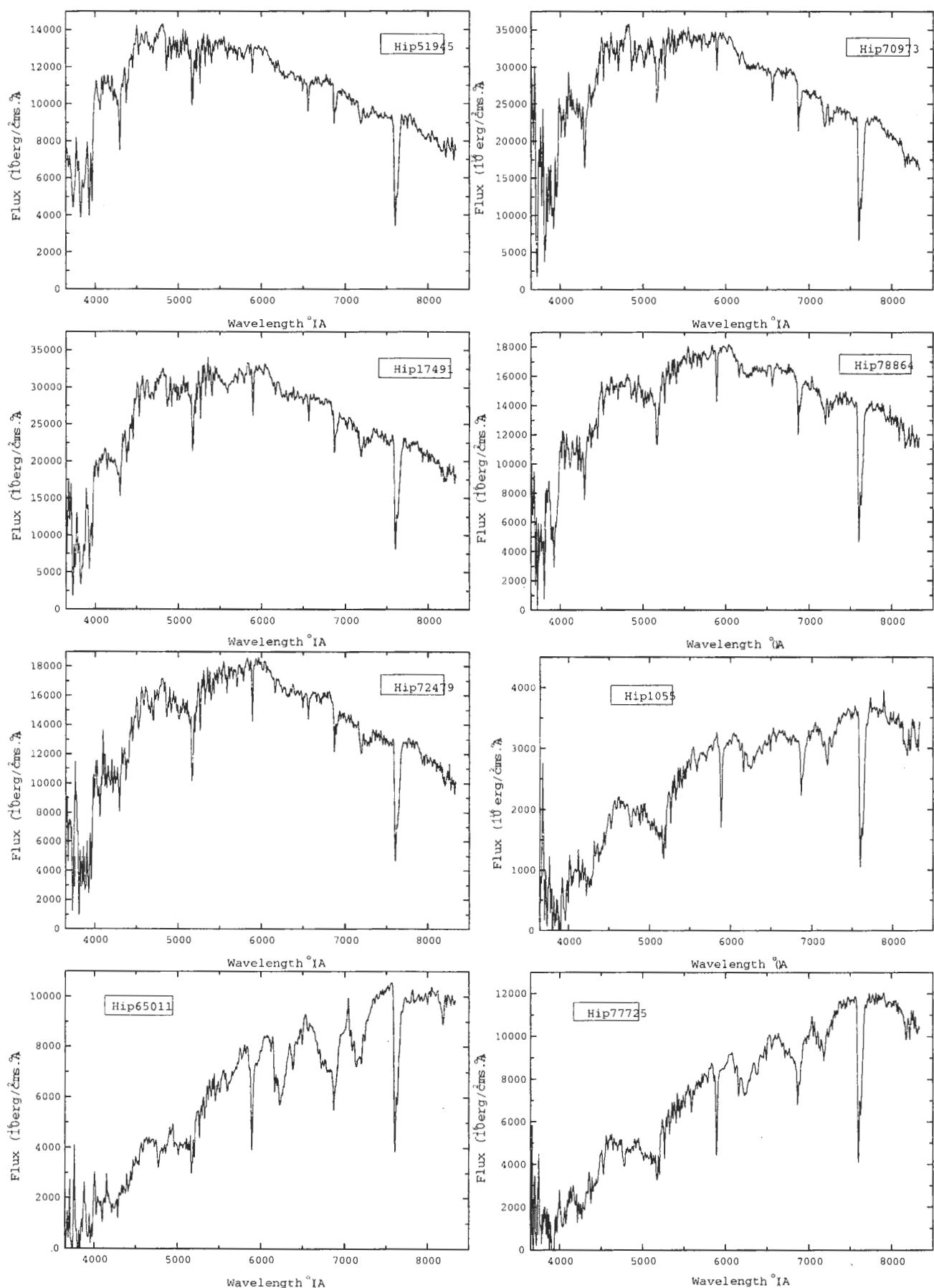


Figure 5: (Continued).

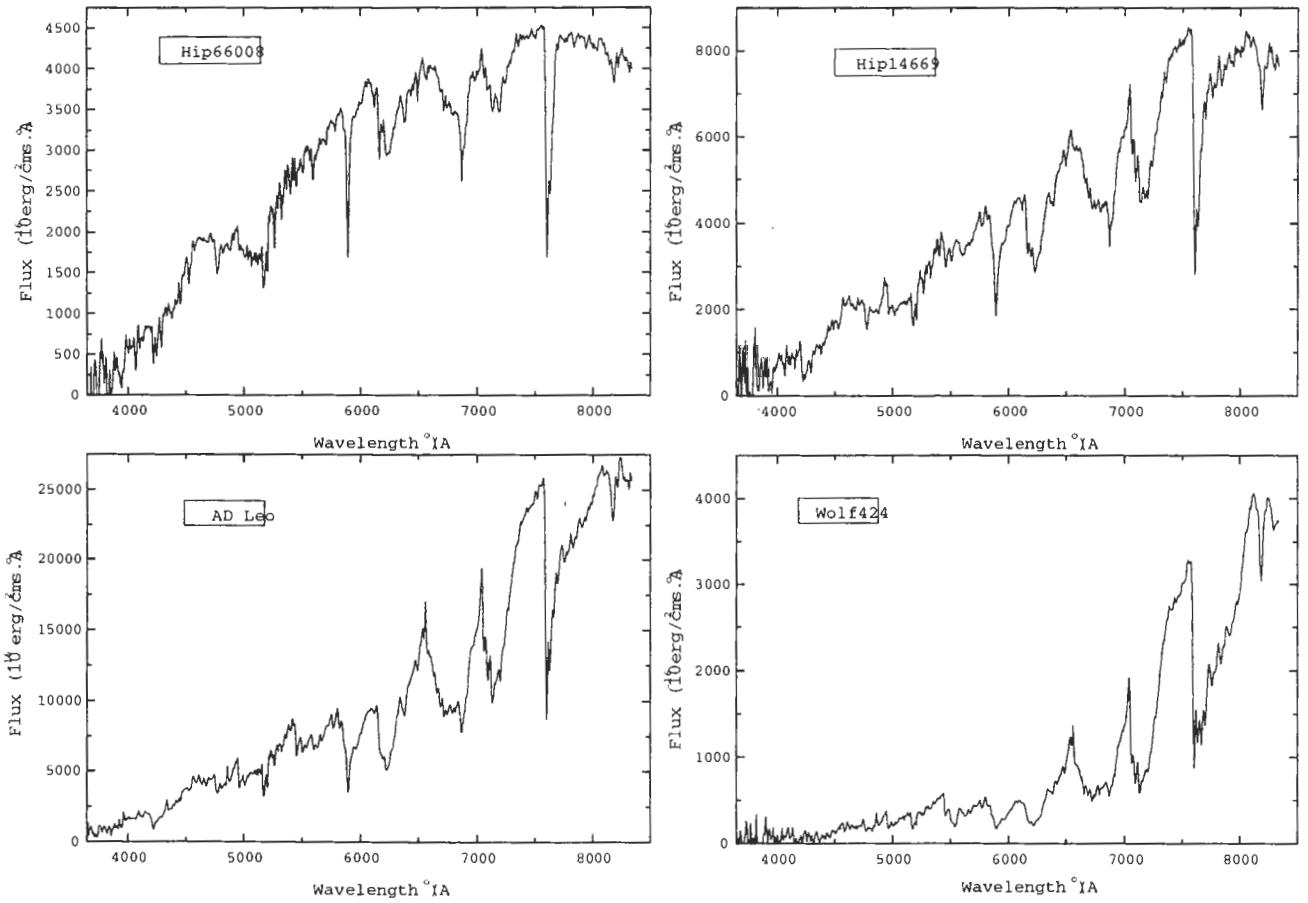
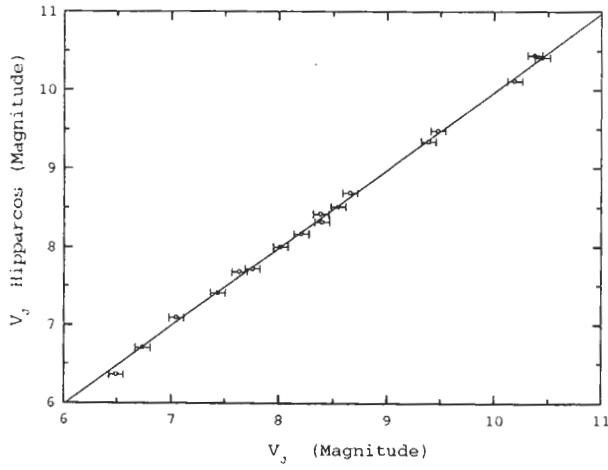
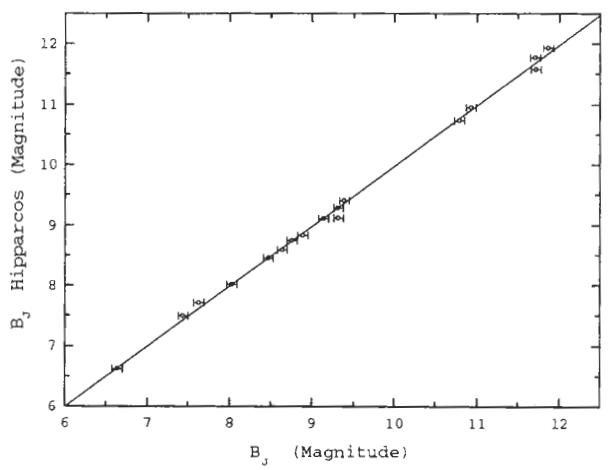


Figure 5: (Continued).

Figure 6: Comparison between the calculated  $V_J$  magnitudes and  $V_J$  magnitudes of Hipparcos and Tycho catalogues.

#### 4. Conclusions

Composite spectral energy distributions of 20 speckle binary stars were measured, from which 17 are G, K, and M dwarfs. The knowledge of the fundamental

Figure 7: Comparison between the calculated  $B_J$  magnitudes and  $B_J$  magnitudes of Hipparcos and Tycho catalogues.

parameters of these late-type binaries is important for the improvement of mass-luminosity relation at the bottom of the main sequence.

The  $BVR$  magnitudes and the  $B - V$  colour indices have been calculated, and the entire spectral

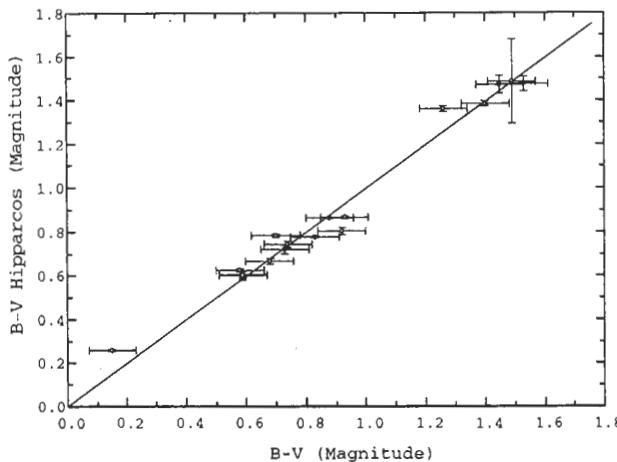


Figure 8: Comparison between the calculated  $(B-V)_J$  and  $(B-V)_H$  of Hipparcos and Tycho catalogues.

types of the pairs have been estimated.

A good agreement has been found between the calculated colour magnitudes and colour indices and those of Hipparcos and Tycho catalogues. Also a comparison of the estimated spectral types with those given by SIMBAD shows a good agreement for 15 stars within the error values of  $B - V$ . Here it is noteworthy that previous papers concerning the quadruple system ADS11061 (Tokovinin, 1995, Al-Wardat, 2002) showed a big difference between the estimated entire

spectral types and those given by SIMBAD.

## References

- Al-Wardat M.A., Balega Yu.Yu., Weigelt G., Vlasyuk V.V., Leushin V.V., Pluzhnik E.A., 2002, (in preparation)
- Al-Wardat M.A., 2002, Bull. Spec. Astrophys. Obs., **53**, 51
- Balega I.I., Balega Yu.Yu., Hofmann K. H., Maksimov A. F., Pluzhnik E. A., Schertl D., Shkhagoshova Z.U., Weigelt G., 2002, Astron. Astrophys., **385**, 87
- Bessel M.S., 1990, Publ. Astr. Soc. Pacific, **102**, 1181
- ESA, 1997, The Hipparcos and Tycho Catalogues, European Space Agency, **1**, 57
- FitzGerald M., 1970, Astron. Astrophys., **4**, 234
- Hamuy M., Suntzeff N.B., Heathcote S.R., Walker A.R., Gigoux P., Phillips M.M., 1992, Publ. Astr. Soc. Pacific, **104**, 533
- Hamuy M., Suntzeff N.B., Heathcote S.R., Walker A.R., Gigoux P., Phillips M.M., 1994, Publ. Astr. Soc. Pacific, **106**, 566
- Hamuy M., Philip A. P., Maza J., et al., 2001, Astrophys. J., **558**, 615
- Hayes D.S., 1985, in: Calibration of Fundamental Stellar Quantities, IAU Symp. 111, eds.: D.S. Hayes, L.E. Pasinetti, and A.G. Davis Philip (Dordrecht: Reidel), 225
- Massey P., Strobel K., Barnes J.V., Anderson E., 1988, Astrophys. J., **328**, 315
- Oke J.B., 1990, Astron. J., **99**, 1621
- Tokovinin A.A., 1995, Pis'ma Astron. Zh., **21**, 250

Table 1: List of the stars with  $B_J$ ,  $V_J$ ,  $R_C$ ,  $(B - V)_J$ , and spectral type results of this work.

Hip (1)	Star Name HD (2)	Other Identifications (3)	$\alpha_{2000}$ ( $h$ )( $m$ )( $s$ ) (4)	$\delta_{2000}$ ( $^{\circ}$ )( $'$ )( $''$ ) (5)	Times of Obs. JD 2452300+ (6)	$B_J$ mag $\pm 0.06^*$ (7)	$V_J$ mag $\pm 0.06^*$ (8)	$R_C$ mag $\pm 0.07^*$ (9)	$(B - V)_J$ $\pm 0.08^*$ (10)	Sp. Type This Work (11)	Sp. Type SHIMBAD (12)
689	375		00 08 28.47	+34 56 04.37	10.171, 10.201	8.02	7.42	7.17	0.6	G0	F8
1055	-		00 13 09.17	+20 22 56.75	10.181, 10.207	11.70	10.42	9.64	1.28	K7	M0
4809	6009		01 01 43.58	+25 17 31.98	3.195, 10.190, 10.212	7.43	6.73	6.32	0.70	G6	G8IV
11253	14874		02 24 51.15	+30 38 48.28	3.208, 10.275, 10.279	8.88	8.20	7.80	0.68	G5	G0V
11352	15013		02 26 09.59	+34 28 10.03	3.213, 10.279, 10.301	8.75	8.00	7.60	0.75	G8	G5
12552	16656	COU 1511	02 41 28.88	+40 52 50.83	3.250, 10.283, 10.305	9.13	8.54	8.26	0.59	G0	G0
14230	18490		03 03 28.66	+23 03 41.33	10.285, 10.309	7.62	7.04	6.68	0.58	F9	G0
14669	-	GJ 125	03 09 30.79	+45 43 57.87	10.313, 10.348	11.71	10.16	9.13	1.54	M4	M2
17491	23140	GJ 150.2	03 44 48.83	+46 02 09.04	10.319, 10.358	8.63	7.74	7.31	0.88	K1	K2
-	25811	CHA 13	04 06 16.41	+19 52 28.58	10.322, 10.366	9.03	8.63	8.44	0.4	F3	F0
-	AD Leo		10 19 36.27	+19 52 11.9	10.422	10.87	9.34	8.20	1.53	M4.5	M4.5V
51945	91877	BAG 3	10 36 43.07	+11 01 18.96	10.433, 10.430	9.38	8.64	8.24	0.74	G8	K0
-	Wolff424		12 33 17	+09 01 20	10.495	14.25	12.50	10.96	1.74	M7	M5Ve
64838	115488	FIN 350	13 17 29.85	-00 40 33.83	10.524, 10.591	6.63	6.47	6.39	0.15	A5	F0V
65011	-	GL 507A	13 19 33.59	+35 06 36.56	10.508, 10.577	10.92	9.46	8.61	1.47	M0	M0
66008	-	WOR 24	13 31 58.25	+31 08 04.56	10.515, 10.583	11.86	10.36	9.47	1.51	M4	M0V
70973	-	RST 4529	14 31 00.62	-05 48 08.47	10.597, 10.633	8.46	7.62	7.21	0.84	K0	G5
72479	-	ADS 9397	14 49 20	+10 14 00	10.607, 10.637	9.31	8.37	7.88	0.95	K3	K2V
77725	-	GL 600	15 52 08.24	+10 52 28.11	10.613, 10.644	10.78	9.37	8.46	1.41	M0.5	M2
78864	144515	NQ Ser	16 05 53.41	+10 41 06.01	10.618	9.31	8.38	7.86	0.92	K2	G8IV

\* The error values for most of the stars are better than these values.























Table 2: Flux ( $E-16 \text{ erg/cm}^2 \cdot \text{s} \cdot \text{\AA}$ ) (continued)

$\lambda$	A	HIP	HIP	HIP																	
8130	13680	10556	48860	11253	11352	12552	14230	14669	17491	25811	AD	51945	424	64838	65011	66008	70973	72479	77725	78864	
8135	13467	30685	48860	11601	14879	63778	32445	8139	20290	5396	26198	7641	4033	34467	10248	4166	18948	11105	11162	12593	
8142	13693	47716	11997	14013	14016	6453	32277	8123	19932	5270	26217	7820	3997	33265	9887	4197	18729	11044	11248	12282	
8148	14323	49076	48194	11832	14016	6511	31961	8080	19657	4947	26137	7692	33393	9796	4136	18408	11052	11122	12396		
8154	13234	30688	48888	11201	13969	6074	29467	7804	19090	5246	25455	7520	3882	34429	9739	4035	18473	10845	10980	12853	
8160	13392	30556	48888	11086	13415	6058	28516	7670	19181	5103	25015	7421	3846	33884	9815	4068	17867	11500	10681	12167	
8166	13894	30111	48877	10708	13120	5910	29066	7676	17975	4769	23920	7511	3773	32735	9550	4110	16516	10782	10697	11086	
8172	14220	29779	47376	10734	13237	6004	29467	7636	19152	4896	23020	7423	3560	31438	9478	4026	17588	11184	10287	11530	
8178	13541	47003	29660	11340	13875	6113	29426	7417	19232	4748	22779	7469	3280	32526	9378	3932	17924	10996	9866	11281	
8184	12877	2883	48003	11209	14308	5993	29968	7052	18837	4919	22987	7432	3099	32043	9101	3863	18030	10294	9912	11272	
8190	13608	30303	48627	11422	14028	5844	29126	7654	18155	5182	23654	7573	3044	30759	8864	3825	16347	10308	10196	11534	
8196	14025	3161	49679	11504	14011	5738	29154	6664	17146	5098	24913	7906	3181	33201	9004	3915	17597	9980	10531	11651	
8202	13630	31116	49656	11755	13755	5948	29960	6991	17578	5179	25913	7987	3434	33672	9330	4017	18375	10512	10848	11686	
8208	12627	3173	48884	11422	13682	5914	29407	7364	17973	5200	25948	7627	3647	35425	9712	4137	18225	9815	10664	11999	
8214	12422	29559	47085	10655	12943	6251	29382	7664	17929	5164	25500	7268	3751	33661	10049	4198	17490	10132	10544	12513	
8220	12626	29669	6453	10226	12611	6274	29721	7636	17248	4891	25183	7116	32133	10071	4097	16539	9952	10080	11850		
8226	13256	3295	46395	10571	11549	6180	29435	7616	17397	5104	25637	7380	3858	30892	9791	4013	17539	10277	9866	11227	
8232	14027	3243	48985	10840	11604	6044	29846	7591	17944	5316	26750	7762	3874	32717	9842	4110	17352	10476	10608	11266	
8238	13681	3357	49757	11001	11971	6102	30372	7642	19303	5139	27188	8043	4003	32717	9842	4154	17515	10765	10580	11491	
8244	13381	33116	49656	11422	13682	5914	29407	7364	17973	5200	25948	7627	3647	35425	9712	4137	18225	9815	10664	11999	
8250	13676	34884	47377	12397	14013	63778	32445	8139	20290	5396	26198	7641	4033	34467	10248	4166	18948	11105	11162	12593	
8256	13630	3332	49654	11247	13707	6235	29444	8191	19131	5378	26827	7726	4002	33146	9887	4239	17988	10218	10776	12320	
8262	13525	3421	48215	11140	14605	6154	29721	8071	19703	5036	26365	7655	3946	31560	10005	4227	17823	11197	112049	12049	
8268	13297	3344	49774	11779	14780	5881	30513	7977	18749	4975	25762	7569	3916	31736	9901	4148	17797	10389	11212	12049	
8274	12703	3349	49407	11528	14115	5933	30325	8121	18244	5118	25678	7542	3888	32661	9880	4149	17357	10245	10495	11612	
8280	11478	3189	50129	10984	13008	6101	29508	7978	18174	4892	25928	7437	3800	31919	9734	4116	17698	9554	10362	11192	
8286	12619	3132	49277	10741	12397	6173	30048	7827	18322	5179	27298	8179	4016	33157	10050	4164	17597	10907	11242	11974	
8292	13232	3059	49652	12004	12397	6274	29632	8084	18176	5224	27245	7885	4008	32275	10087	4215	1752	10327	10810	12711	
8298	13207	3013	48215	12349	12889	6166	28408	7675	18608	4824	25650	8169	3646	31906	10083	4124	17259	9900	10501	11711	
8304	12736	3135	47503	12366	1332	6009	27752	7550	18841	4752	25724	7665	3669	32096	9937	4104	16578	9949	10217	12042	
8310	12044	47820	11577	13610	5590	28713	7628	19480	4426	25666	7180	3700	30796	9882	4037	17557	10273	10398	11928		
8316	12616	2967	48494	11361	13027	5540	30046	7890	17928	18172	4672	25060	6961	9776	4001	17118	10050	10107	11232		
8322	13056	3312	48442	10819	12836	5705	30632	7718	18311	4772	26235	7801	3755	31454	9693	4082	16804	10473	10344	11146	
8328	13056	3438	48447	10725	12725	5776	30256	7718	18311	4772	26058	7670	3724	30911	9861	4073	16806	10478	10344	11146	
8334	11891	3378	49608	11296	12404	5578	30332	7644	17765	5245	25637	7494	3742	31195	9782	3990	16426	10345	10345	11752	